

# Long-term multi-frequency maser observations of the intermediate-mass young stellar object G107.298+5.639<sup>★</sup>

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## ABSTRACT

**Context.** Periodic flares of maser emission are thought to be induced either by variations of the seed photon flux in young binary systems or the pump rate regulated by stellar and accretion luminosities.

**Aims.** We seek to study the variability of four maser transitions of three different species in G107.298+5.639 to constrain the dominant mechanism of periodic flares.

**Methods.** Light curves of the 6.7 GHz methanol and 22.2 GHz water vapour maser were obtained with the Toruń 32 m radio telescope over 39 and 34 cycles, respectively. The target was also monitored at the 1.6 GHz hydroxyl transitions with the Nançay radio telescope over 13 cycles. All these maser lines were imaged using VLBI arrays.

**Results.** The study confirms alternating flares of the methanol and water masers with a period of 34.4 d and reveals the synchronised behaviour of the methanol and hydroxyl masers in this source. The observed spatial distribution of the methanol maser cloudlets and the measured time delays of the flares of individual features imply a ring-like structure of radius 240 au and thickness 30 au. Internal proper motions indicate that the velocity of methanol cloudlets is dominated by a disc-wind component of about  $5 \text{ km s}^{-1}$ . The methanol emission detected during only one VLBI observation is located in a region about 550 au from a central star, which also exhibits OH maser flares. The erratic appearance of methanol features can be related to a powering object of relatively low luminosity which, during some variability cycles, can excite molecules only in the nearest part of the disc. A careful analysis of the maser and infrared light curves reveal a strong correlation between the 6.7 GHz line and the infrared flux densities supporting a radiative pumping of the maser.

**Conclusions.** The synchronised behaviour of the hydroxyl 1665/1667 MHz and 6.7 GHz methanol transitions indicates a common pumping mechanism for the periodic flares of G107.298+5.639.

**Key words.** masers – stars: formation – ISM: clouds – radio lines: ISM – instabilities – stars: individual: G107.298+5.639

## 1. Introduction

Maser lines coming from outflows, circumstellar discs, and envelopes are the most characteristic properties of high-mass young stellar objects (HMXOs) for a relatively very short ( $5 \times 10^4$  yr) phase of their evolution (van der Walt 2005; Breen et al. 2010). A small fraction of HMXOs show periodic (24–668 day) variations in the 6.7 GHz  $\text{CH}_3\text{OH}$  maser emission (Goedhart et al. 2003, 2004, 2009, 2014; Szymczak et al. 2011, 2015; Fujisawa et al. 2014; Maswanganey et al. 2015, 2016; Sugiyama et al. 2017) and a few objects also show periodic variations in the  $\text{H}_2\text{CO}$ , OH and  $\text{H}_2\text{O}$  maser lines (Araya et al. 2010; Al-Marzouk et al. 2012; Szymczak et al. 2016).

Hypotheses explaining periodic variations in the methanol flux density are based on changes in the flux of seed photons or radiative pump rate. van der Walt (2011) claim that maser light curves of widely different shape and amplitude in some objects are due to periodic modulation of the background emission, which is amplified by the masering gas. Cyclic changes in the seed photon flux could arise in a colliding-wind binary or an eclipsing binary (van der Walt et al. 2009; van der Walt 2011; Maswanganey et al. 2016). An alternative explanation of the

methanol periodicity is related to the pump rate, which depends on the flux of infrared emission (Cragg et al. 2002). Araya et al. (2010) propose a scenario of periodic heating of the dust due to gas accretion from the circumbinary disc onto protostars or accretion discs. Periodic maser flares could trace the accretion behaviour of circular or eccentric binaries (Muñoz & Lai 2016). The dust temperature can be regulated by rotating spiral shocks in the gap of a circumbinary accretion disc (Parfenov & Sobolev 2014). Flaring of the methanol masers in some objects can be also due to star pulsation driven by the  $\kappa$  mechanism (Inayoshi et al. 2013) or accretion instabilities developed by interactions between the stellar magnetosphere and the accretion disc (D’Angelo & Spruit 2012).

Observations of periodic methanol masers at other maser transitions of well-established pumping schemes can shed more light on the nature of the periodicity. We can expect a correlation of periodic flares for the two maser transitions which share the same pumping mechanisms and anti-correlation if they are different. For the latter case a correlation should occur if the bursts are due to variations of the background emission amplified by the masers. Araya et al. (2010) report correlated quasi-periodic variability between  $\text{H}_2\text{CO}$  (4.8 GHz) and  $\text{CH}_3\text{OH}$  (6.7 GHz) masers in G37.55+0.20. In this object the flares of the excited OH (6.035 GHz) masers were delayed relative to the  $\text{H}_2\text{CO}$  flares, but may be simultaneous with the  $\text{CH}_3\text{OH}$  peaks at corresponding

\* Data used to produce Figs. 1 and 2 are only available at the CDS via anonymous ftp to [cdsarc.u-strasbg.fr](http://cdsarc.u-strasbg.fr) (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/634/A41>

radial velocities (Al-Marzouk et al. 2012). Green et al. (2012) report that the variability of the 1665 MHz and 1667 MHz lines in G12.889+0.489 shows evidence of correlation with the 6.7 GHz line periodic variability. Furthermore, Goedhart et al. (2019) show that OH lines display concurrent flaring with 6.7 and 12 GHz methanol masers in G9.62+0.20E. These tentative findings support a common excitation process of the OH and CH<sub>3</sub>OH maser transitions. Recently, we found a remarkable pattern of anti-correlated alternating flares of the water (22 GHz) and methanol (6.7 GHz) maser transitions in the intermediate-mass young stellar object (IMYSO) G107.298+5.639 (Szymczak et al. 2016). These variations can be due to periodic changes in the infrared radiation field from the central protostar.

In this paper we report on our time-series observations of G107.298+5.639 (G107 hereafter) of maser transitions of three species to confirm the anti-correlated bursts of the methanol and water maser lines (Szymczak et al. 2016) and to examine the behaviour of ground-state OH masers. The numerical models predict that the far-infrared pump together with the effects of line overlap provides the most convincing fit to the OH maser observational data (e.g. Cesaroni & Walmsley 1991; Gray et al. 1991; Cragg et al. 2002). The OH and CH<sub>3</sub>OH masers are radiatively pumped (Cragg et al. 2002) and we report for the first time their correlated variability in this source.

## 2. Observations

### 2.1. Nançay and Torun radio telescopes

The Nançay telescope has a half-power beam-width (HPBW) of 3.5', in the E–W direction and 19' in the N–S direction. The beam efficiency is 0.65, the point source efficiency at zero declination 1.4 K Jy<sup>-1</sup> and the system temperature is about 35 K. All four transitions of the OH ground state were observed simultaneously in both senses of circular polarisation using a 8192 channel auto-correlator spectrometer configured into eight banks of 1024 channels. A bandwidth of 0.39 MHz was used for each bank, yielding a total useful velocity coverage of about  $\pm 30 \text{ km s}^{-1}$ . The spectral bandwidth was centred at  $-15 \text{ km s}^{-1}$  relative to the local standard of rest (LSR). The spectra were taken in the frequency switching mode. The flux density calibration was based on the flux monitoring of W12 and is deemed to be accurate within 10%. On four occasions around the anticipated maxima the 1665 and 1667 MHz lines were observed in full polarisation mode following the procedure described in Szymczak & Gérard (2009). The four Stokes parameters were obtained with a typical rms noise level of about 60 mJy in the I Stokes and the uncertainty of the polarisation measurements was less than 8%. The target was observed in the OH lines on average once a week and twice a month over 12 and 7 cycles, respectively.

The CH<sub>3</sub>OH and H<sub>2</sub>O maser lines were monitored with the Torun 32 m radio telescope in the frequency-switching mode. For methanol the system temperature ranged from 25 to 30 K. The HPBW was 5' and the pointing accuracy was better than  $\sim 25''$  before mid-2016 and  $\sim 10''$  afterwards. The data were calibrated by regular scans on the continuum source 3C123 and daily observations of the methanol maser G32.744–0.076, in which some features are non-variable within 5% (Szymczak et al. 2014). The typical rms noise level in the final methanol spectra of  $0.09 \text{ km s}^{-1}$  resolution was about 0.3 Jy. The uncertainty of the flux density calibration was  $\sim 10\%$ . The 6.7 GHz emission was observed roughly every six hours and two to three days during the active and quiescent phases, respectively, over 39 cycles.

The 22 GHz observations were calibrated by the chopper wheel method and observations of the continuum source DR21, assuming a flux density of 18.8 Jy. The HPBW was 1.7' and the system temperature ranged between 45 and 250 K depending on weather conditions and elevation. The typical rms noise level was 15–20 mK. The observations were corrected for atmospheric absorption and gain elevation effect. The estimated accuracy of the temperature scale was about 20%. The water line was observed every one to two days over 33 cycles.

### 2.2. Very Long Baseline Interferometry

The dates of Very Long Baseline Interferometry (VLBI) experiments were adjusted according to the ephemeris from the 32 m dish monitoring to get data just at flare peaks. The OH 1665 and 1667 MHz maser lines were observed at two epochs: in 2017 November (MJD 58 063) with the EVN (project ES083) and in 2017 June (MJD 57 929) with the VLBA (project BS254A) at phases 0.56 and 0.63, respectively. The methanol maser line at 6.7 MHz was observed with the EVN (projects ES076, ES079A, ES079B, ES079C) at four epochs over the period from 2015 March to 2016 October at phases  $0.50 \pm 0.03$  (Table 1). The H<sub>2</sub>O line was observed with the VLBA in 2017 June (project BS254B) at phase 0.97, i.e. near the water flare peak.

Table 1 presents an overview of all the observed maser transitions, including the arrays, dates, spectral resolutions, synthesised beam sizes, measured brightness of phase calibrator, and the rms image noise. For all the experiments we used a phase-referencing observation scheme with J2223+6249 as a phase calibrator with a switching cycle of 195 s+105 s (maser + phase calibrator). 3C454.3 was used as a delay and bandpass calibrator. The typical total time on source was  $\sim 5$  h. Data were correlated with 1 s integration time with the exception of experiment ES076, for which the integration time was 0.25 s.

The data reduction followed standard procedures for calibration of spectral line observations using the AIPS package. For EVN the initial step of calibration was conducted on continuum pass and the solutions were applied to the spectral data sets, while VLBA calibrations were performed directly on spectral pass. The target was then self-calibrated on a channel representing the centre of the brightest spectral feature and the obtained phase and amplitude corrections were applied to the whole spectral data cube. For the 1665 MHz transition the emission in the reference velocity component was faint and spatially resolved out in long baselines ( $> 1000$  km), so that the visibility calibration solutions were obtained from the data of the shorter baselines. Stations providing longer baselines or flagged owing to technical failures are denoted in italics in Table 1. No emission at 1667 MHz was detected with both arrays.

Channel maps were searched for emission and the parameters of the detected maser emission were derived following the method presented in Bartkiewicz et al. (2016) and Olech et al. (2019). Point-like emission in each spectral channel is defined as spot. The position, intensity, and flux of each spot have been derived by fitting a 2D elliptical Gaussian. A group of spots coinciding within less than half of the synthesised beam size and present in adjacent channels was termed cloudlet (e.g. Sanna et al. 2017); the average flux-weighted position was calculated. Cloudlets often formed smooth Gaussian spectral profiles and their parameters such as LSR peak velocity ( $V_p$ ), full width at half maximum (FWHM), peak brightness ( $S_p$ ), and fitted brightness ( $S_f$ ) were derived (Sect. 3.5). The absolute positional accuracy of the maser spots estimated using formula described in

**Table 1.** Details on VLBI observations.

Telescope	Date MJD	Frequency (MHz)	Spectral res. (km s <sup>-1</sup> )	Synthesised beam ( $\alpha \times \delta$ ; PA) (mas $\times$ mas; °)	$1\sigma_{\text{rms}}$ per channel (mJy beam <sup>-1</sup> )	J2223+6249 flux (mJy beam <sup>-1</sup> )
EVN <sub>(Jb,Wb,Ef,Mc,O8,T6,Ur,Tr,Sv,Zc,Bd,Ir)</sub>	58 063	1665.402	0.35	30 $\times$ 23; 78.00	6	108
		1667.359	0.35	30 $\times$ 23; 57.65	6	
VLBA <sub>(BR,FD,HN,KP,LA,MK,NL,OV,PT,SC)</sub>	57 929	1665.402	0.09	35 $\times$ 33; 39.07	12	126
		1667.359	0.09	35 $\times$ 33; 10.93	12	
EVN <sub>(Jb,Ef,Mc,O8,Tr,Nt,Wb,Ys,Sr)</sub>	57 097	6668.519	0.09	5.8 $\times$ 3.7; -78.20	1.5	147
EVN <sub>(Jb,Ef,Mc,O8,Tr,Nt,Wb,Ys,Sr)</sub>	57 444	6668.519	0.09	5.2 $\times$ 3.2; 61.38	3.0	189
EVN <sub>(Jb,Ef,Mc,O8,Tr,Nt,Wb,Ys,Sr,T6,Ir)</sub>	57 544	6668.519	0.09	4.2 $\times$ 2.8; 76.25	2.5	172
EVN <sub>(Jb,Ef,Mc,O8,Tr,Nt,Wb,Ys,T6)</sub>	57 681	6668.519	0.09	4.2 $\times$ 3.1; 79.75	5.0	171
VLBA <sub>(BR,FD,HN,KP,LA,MK,NL,OV,PT,SC)</sub>	57 906	22 235.080	0.11	0.6 $\times$ 0.5; -1.10	5.0	104

**Notes.** No data were used from italicised stations.

Bartkiewicz et al. (2009) are 18, 3, and 3 mas at 1.6, 6.7, and 22 GHz, respectively.

To measure the proper motion of the methanol masers we used following procedure: the data at all epochs were phase referenced and calibrated using the AIPS package in an identical way. We used the first epoch as a reference to find the persistent maser clouddlets at the other epochs. Any systematic difference in the clouddlet position between the observations due to residual phase referencing errors, annual parallax, and source motion were removed by calculating the mean positions of clouddlet distributions, i.e. the barycentres of the masers. If after those correction any clouddlet drastically changed its position at any epoch, it was deemed as not persistent and removed from the analysis. Finally the proper motion estimation of individual clouddlets was performed with linear fit of right ascension (RA) and Dec displacement over all epochs.

### 3. Results

#### 3.1. Maser light curves

The integrated flux density time series for the four maser transitions presented in Fig. 1 show clear periodic variations. The amplitude of the individual features and shape of spectra displayed significant variability from cycle to cycle. These complex variations are best seen in the plots of the flux density as a function of time and velocity (Figs. A.1 and A.2).

Following the approach used in Olech et al. (2019) we calculated the period and analysed the flare profiles for all the transitions. An asymmetric power function of the form  $S(t) = A s(t) + C$  was fitted to the data. In this equation,  $A$  and  $C$  are constants and  $s(t) = [b \cos(\omega t + \phi)/(1 - f \sin(\omega t + \phi))] + a$ ,  $b$  is the amplitude relative to mean value  $a$ ,  $\omega = 2\pi/P$ , where  $P$  is the period,  $\phi$  is the phase, and  $f$  is the asymmetry parameter (Szymczak et al. 2011). The resulting fits were used to assess the time of flare peak, the timescale of variability equals to the FWHM of the flare, the relative amplitude ( $R$ ), and the ratio of the rise time to the decay time ( $R_{\text{rd}}$ ). Table 2 summarises their average values for all the spectral features. The active phase is the shortest for the 6.7 GHz line; the average FWHM is only 4.30 d and the flare profile is fairly symmetric with a mean  $R_{\text{rd}} = 0.92$ . The flare profile at 22 GHz has an average FWHM of 16.34 d and the rise time is longer than the decay time ( $R_{\text{rd}} = 1.21$ ). Both OH transitions have similar FWHM of  $\sim 10$  d but the 1665 MHz line has a steeper flare profile ( $R_{\text{rd}} = 0.61$ ) than the 1667 MHz line ( $R_{\text{rd}} = 0.93$ ).

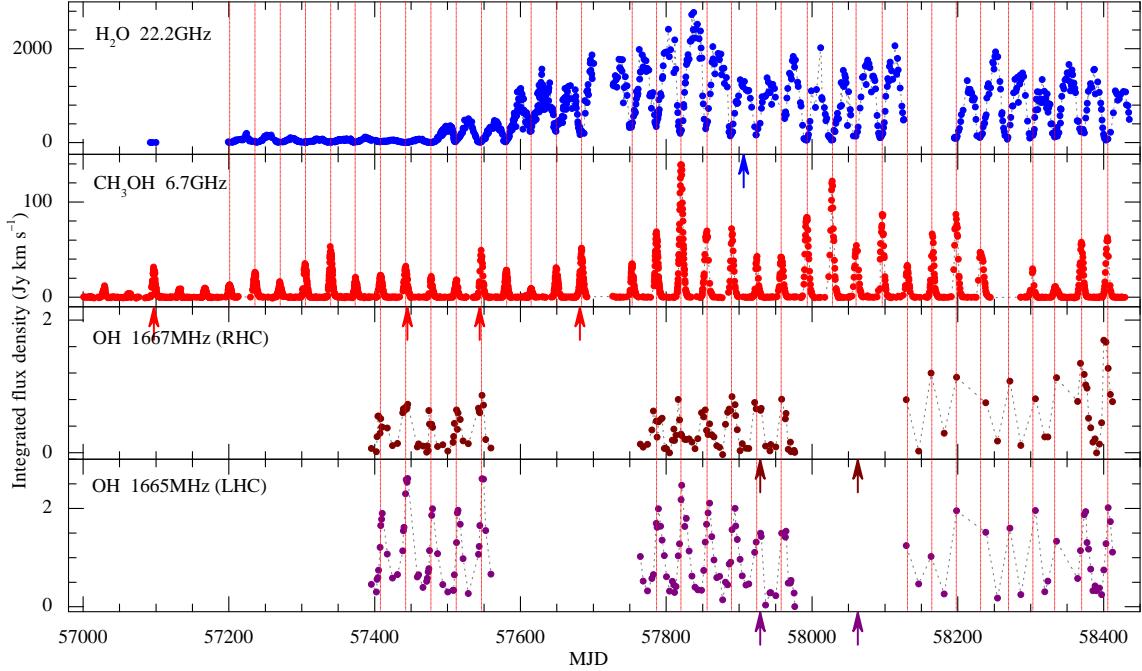
The time series of the velocity-integrated flux densities were normalised and phase-folded to show differences in the flare profiles and overall time delays between the peaks of observed transitions (Fig. 2). The anti-correlated behaviour of the water vapour and methanol masers is clearly visible; the minimum of water flux density coincides exactly with the maximum of methanol emission. Since the essential mechanisms of pumping H<sub>2</sub>O and CH<sub>3</sub>OH population inversions are collisional and radiative, respectively (Elitzur et al. 1989; Cragg et al. 2002), the observed anti-correlation rules out the hypothesis that the maser flares in the target are related to variations of seed photon flux caused by periodic changes in the free–free background emission in a system of colliding-wind binary. The methanol and OH 1665 MHz light curves are quite a bit different, as the asymmetric model fits in Fig. 2 suggest; the  $R_{\text{rd}}$  ratio is smaller at 1665 MHz, where the peak is flat-topped followed by a slow decay as often seen in OH masers associated with the Mira variables (e.g. Etoka et al. 2017). The phased light curve of the 1667 MHz emission is less reliable owing to a large scatter of the data points and indicates no time lag relative to the methanol flare.

#### 3.2. Periodicity estimates

A Lomb-Scargle (L-S) periodogram as implemented in the VARTOOLS package (Hartman & Bakos 2016) was used to search for any periodicity in the time series of the velocity-integrated flux densities in each of the four transitions shown in Fig. 1. We followed the procedure described in detail by Olech et al. (2019). The L-S periodograms are shown in Fig. 3. The periods with the highest power for the water and hydroxyl 1665 and 1667 MHz masers are  $34.35 \pm 0.51$ ,  $34.36 \pm 0.48$  and  $34.35 \pm 0.49$  d, respectively. These values are the same, within the error range, as for the methanol maser of  $34.41 \pm 0.72$  d. The periodogram of the methanol time series clearly shows second and third harmonic signals of the derived period at 17.2 and 11.5 d and the same responses are present for the 1665 MHz signal as well. The peak corresponding to the methanol periodicity clearly splits as a consequence of slight variations in the flare profile parameters from cycle to cycle and varied contribution of individual features to the velocity-integrated flux density as Table 2 depicts.

#### 3.3. Time delay

We calculated the discrete correlation function (DCF; Edelson & Krolik 1988) to determine time delays between the flare peaks



**Fig. 1.** Time series of integrated flux densities in the  $I$  Stokes for water vapour and methanol maser transitions and in the RHC and LHC polarisation for hydroxyl maser transitions of G107. The vertical dotted lines indicate the methanol peaks. The arrows below the time axes denote the dates of VLBI observations (Table 1).

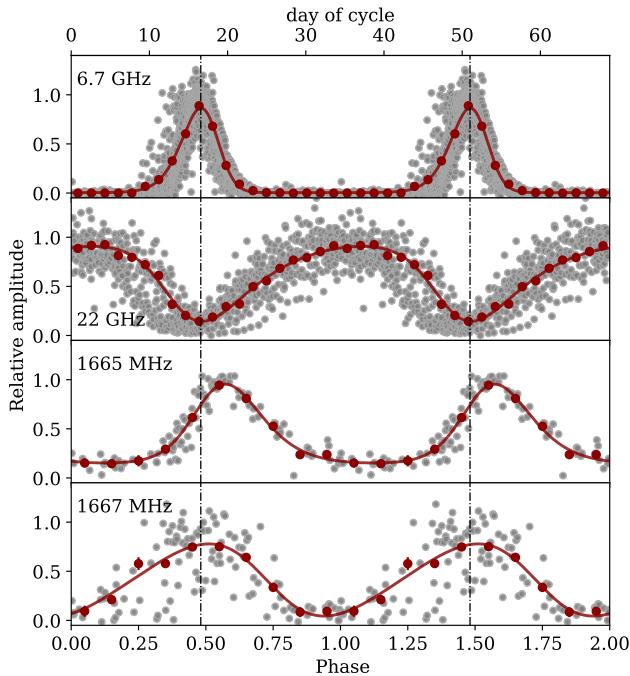
**Table 2.** Characteristics of the maser flares of persistent features inferred from the long-term monitoring.

Transition	$V$ ( $\text{km s}^{-1}$ )	$S$ (Jy)	$R$	$FWHM$ (days)	$R_{\text{rd}}$	$\tau$ (days)
$\text{CH}_3\text{OH}$ 6.7 GHz	-16.78	0.61	>6.00	2.99 (0.82)	1.22 (0.39)	-2.09 (0.51)
	-16.21	0.26	>2.10	3.10 (0.68)	1.16 (0.32)	-1.94 (0.52)
	-15.60	0.97	>7.34	5.03 (1.05)	0.80 (0.23)	3.67 (0.28)
	-14.85	0.17	>0.94	5.03 (1.03)	0.80 (0.20)	1.43 (0.39)
	-13.75	0.17	>1.38	4.08 (0.62)	0.88 (0.16)	2.95 (0.25)
	-11.03	1.18	>4.70	5.82 (1.33)	1.0 (0.43)	1.58 (0.19)
	-9.23	4.42	>31.29	3.78 (0.71)	0.89 (0.24)	-0.11 (0.22)
	-8.62	4.16	>19.77	4.73 (1.20)	0.70 (0.26)	-1.93 (0.27)
	-7.43	16.60	>105.81	4.10 (0.87)	0.84 (0.22)	0
OH 1.665 GHz	-12.22	1.65	>2.39	9.46 (0.60)	0.67 (0.35)	1.90 (0.53)
	-10.03	0.44	>5.94	10.83 (1.06)	0.54 (0.15)	2.91 (0.44)
OH 1.667 GHz	-14.00	0.65	>1.41	9.87 (1.94)	0.77 (0.38)	-2.63 (0.53)
	-9.62	0.32	>1.25	11.10 (3.83)	1.10 (0.50)	-1.02 (0.90)
$\text{H}_2\text{O}$ 22 GHz	-18.83	42.00	22.49 (18.01)	15.61 (2.32)	1.29 (0.48)	-0.19 (0.41)
	-16.40	218.60	146.51 (64.10)	17.62 (1.49)	1.19 (0.44)	0
	-7.85	36.59	18.70 (26.68)	15.79 (1.86)	1.15 (0.49)	-0.50 (0.64)

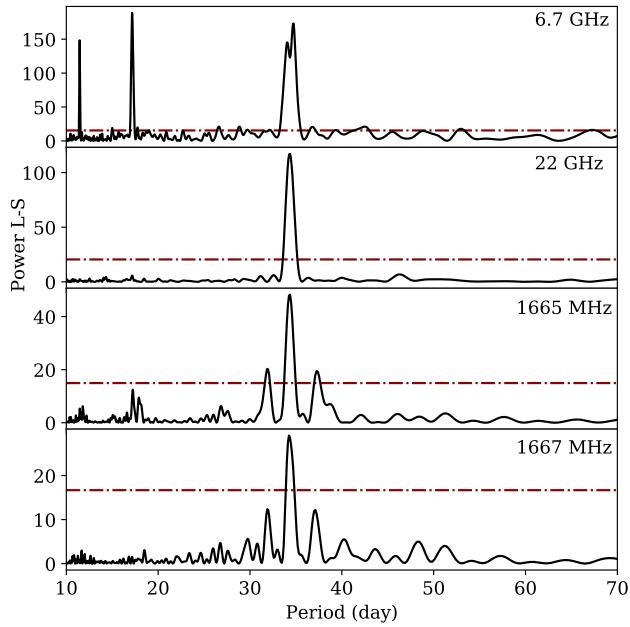
**Notes.** The quantity  $V$  is the peak velocity of the feature. Average values are given for the flux density ( $S$ ), relative amplitude ( $R$ ), timescale of flare ( $FWHM$ ), ratio of the rise time to the decay time ( $R_{\text{rd}}$ ), and time delay ( $\tau$ ). Delays for the methanol and water maser lines were calculated relative to the features  $-7.43$  and  $-16.40 \text{ km s}^{-1}$ , respectively.

of persistent spectral features. The  $-7.43 \text{ km s}^{-1}$  feature of the 6.7 GHz spectrum was chosen as a reference time series for each feature. The time delay was estimated by fitting a quadratic function to the DCF maximum as demonstrated in Olech et al. (2019). The results are presented in the last column of Table 2 for the methanol and hydroxyl transitions. Thanks to high cadence observations and the relatively narrow profile of the methanol

flare the estimated delays have an uncertainty of a fraction of a day. The time delays of the methanol features range from  $-2.1$  to  $+3.7$  d, which well fit the line-of-sight effect for a typical size of maser source of about 1000 au (Bartkiewicz et al. 2009). Time delays between the water maser features are less than 0.5 d and suffer large uncertainties resulting from the strong blending effect of spectral components and widespread drifts in velocity.



**Fig. 2.** Phased and normalised light curves of the integrated flux in each observed transition (grey points). Red points show binned light curves, with  $0.05P$  and  $0.1P$  bins for the 6.7, 22 GHz and both OH lines, respectively. The red line represents the best fit of asymmetric periodic power function. All the light curves are folded modulo 34.43 d, the variability period of the methanol emission, and the dashed vertical lines show the peaks of the 6.7 GHz flare.



**Fig. 3.** Lomb-Scargle periodograms (black lines) calculated for the integrated flux time series in each of the observed transitions. The red dashed line represents a 0.01% false alarm probability level.

#### 3.4. Polarised OH emission

Figure A.3 shows the full polarisation spectra of the OH 1665 and 1667 MHz lines, respectively. Our convention for the calculation, scaling, and displaying of polarisation parameters is the same as in Szymczak & Gérard (2009).

The 1665 MHz emission ranges from  $-12.5$  to  $-7.4 \text{ km s}^{-1}$  with a peak flux density of 2.1 Jy at  $-12.2 \text{ km s}^{-1}$  on MJD 57 821, i.e. very close to the methanol maser maximum (Fig. A.3). Left-hand circular (LHC) polarisation emission dominates and the degree of circular polarisation ranges from 55% to 95%. The main feature at  $-12.2 \text{ km s}^{-1}$  is also linearly polarised up to 18% and the polarisation angle is  $-75 \pm 10^\circ$ . For all OH polarisation observations carried out very near to the methanol peaks at four consecutive cycles (from MJD 57 787 to 57 891) the polarisation angle remains constant within errors. The emission of the main 1665 MHz feature was seen over the entire cycle of variability (Fig. A.2). The 1667 MHz spectrum is composed of two weak ( $\lesssim 0.5$  Jy) features near  $-14$  and  $-9.5 \text{ km s}^{-1}$ . Only the right-hand circular (RHC) polarisation emission is detected. The 1667 MHz flux density is above our detection limit of 240 mJy only during the methanol flares.

In general, the OH maser emission does not coincide in velocity with the methanol and water masers with the exception of a faint OH 1665 MHz feature near  $-7.4 \text{ km s}^{-1}$ , which appears at the same velocity as the strongest methanol feature. The velocity-integrated flux density of the 1665 MHz emission is a factor of 3 higher than that at 1667 MHz.

#### 3.5. Maser spatial distribution

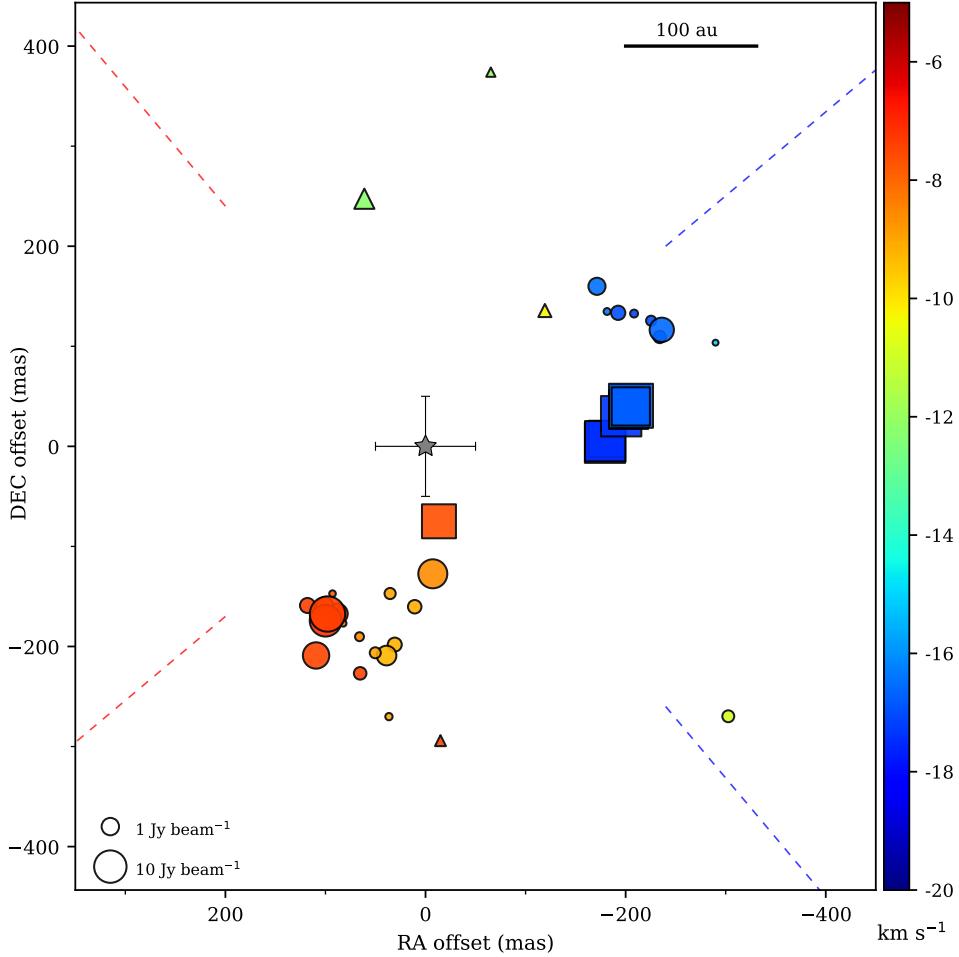
Figure 4 shows the positions and radial velocities of the 22 GHz water, the 6.7 GHz methanol, and the 1665 MHz hydroxyl masers observed very close to the maser peaks but at different cycles spanning less than 1.1 yr. Table A.1 presents the measured and derived maser properties: relative position, peak velocity, velocity FWHM of the cloudlet, brightness of the strongest spot of each cloudlet, and fitted peak brightness of each cloudlet.

During the fourth epoch (MJD 57 681) the methanol emission emerged from two groups of cloudlets  $\sim 400$  mas apart along the NW–SE direction with position angle of  $-43^\circ$  and a lonely cloudlet located in the SW part of the structure (Fig. 4). There is a clear segregation in radial velocity of the NW ( $-17.1$  to  $-15.0 \text{ km s}^{-1}$ ) and SE ( $-9.3$  to  $-7.1 \text{ km s}^{-1}$ ) groups of methanol masers that agrees well with the velocities of blue- and red-shifted lobes of outflows observed in the CO lines (Palau et al. 2011).

The water maser emission was detected in two clusters placed within  $\sim 200$  mas of each other just between the two main groups of methanol masers (Fig. 4). These water maser clusters differ in radial velocity in a similar way to those of methanol masers. It is interesting that the relative positions and radial velocities of these water maser clusters are similar to those seen ten years ago (Hirota et al. 2008). However, the morphology of the western cluster does not resemble that observed by Hirota et al. (2008), who interpreted it as an expanding shell. Moreover, at the epoch of our observation (MJD 57 901) no water maser was detected in the range from  $-3.6$  to  $-3.1 \text{ km s}^{-1}$  and from  $-14.3$  to  $-9.7 \text{ km s}^{-1}$  reported in Hirota et al. (2008).

The OH masers were successfully mapped only at 1665 MHz. Three blue-shifted and one red-shifted cloudlet are detected at the north and south parts of the structure, respectively (Fig. 4). Their distribution seen with different interferometers remains stable at two epochs spanning 4.4 months. We note that the OH emission is completely resolved out at long baselines ( $> 1000$  km) and the positions of cloudlets in Fig. 4 denote the centroids of more extended structures of low brightness.

The overall structure of the red-shifted ( $> -11.3 \text{ km s}^{-1}$ ) methanol maser remained the same at all four epochs. During the second epoch (MJD 57 444) the blue-shifted emission had



**Fig. 4.** Composite map of three maser transitions in G107. The OH 1665 MHz (triangles),  $\text{CH}_3\text{OH}$  6.7 GHz (circles), and  $\text{H}_2\text{O}$  22 GHz (squares) cloudlets are shown; the observations were carried out on MJD 58 063, 57 681 and 57 929, respectively. The symbol size is proportional to the logarithm of maser brightness of the cloudlets and its colour denotes the velocity according to the colour bar. The systematic errors in the relative position between different transitions are estimated to be less than 18 mas. The grey star symbol with the marked positional uncertainty represents the position of the 1.3 mm continuum peak at  $\text{RA}(\text{J}2000)=22^{\text{h}}21^{\text{m}}26\text{:}7730$  and  $\text{Dec}(\text{J}2000)=63^{\circ}51'37\text{.}657$  (Palau et al. 2013) likely indicating the location of IMYSO. The blue and red dashed lines denote the directions of two large-scale molecular outflows (Palau et al. 2011). The horizontal bar indicates the linear scale for the assumed distance of 0.76 kpc (Hirota et al. 2008).

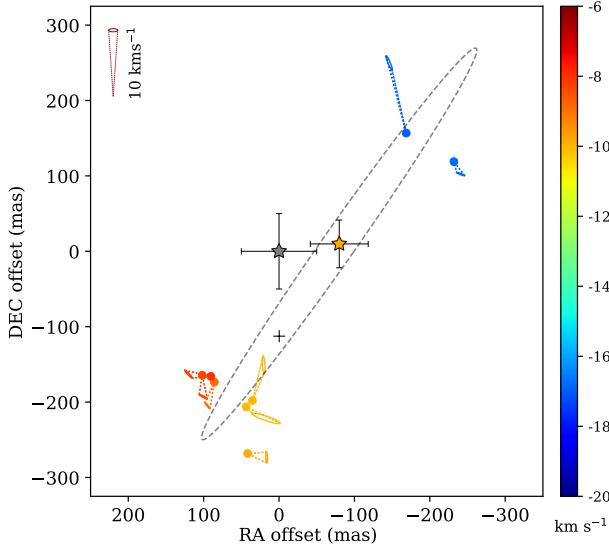
an entirely different structure than at the other three epochs (Fig. A.4); the emission from the NW location did not appear at all, but four new cloudlets at slightly different velocities emerged in the north about 500–600 mas from the brightest red-shifted clouddlet near  $-7.4 \text{ km s}^{-1}$ . The ratio of the peak flares measured by the velocity-integrated flux density at the second and fourth epochs of VLBI observations was as 7:11. This ratio for the first and fourth VLBI epochs was 8:11, while the overall structure remained unchanged. All four VLBI observations were taken within less than 0.09 of phase around the flare maxima. This corresponds to 3 d and is a significant fraction ( $\sim 30\%$ ) of the FWHM of feature flare. Thus the observed methanol morphology is highly dependent on the phase of the flare and the specific variability of the features for a given cycle as shown in Fig. A.5.

### 3.6. Proper motion

In order to establish the correspondence of 6.7 GHz maser cloudlets over the epochs, first we employed the two criteria proposed by Sanna et al. (2010): (i) persistence of the relative distribution of a group of cloudlets and (ii) assumption of uniform motions. The second epoch data are excluded from this study since they did not meet the first criterion; the NW group of cloudlets was not seen while a new emission appeared about 600 mas north of the persistent south structure. Thus for the purpose of proper motion measurements we used only the data from the remaining three epochs. Following the procedure described in Sect. 2.2, we identified eight persistent cloudlets. Linear fits

of displacements of the cloudlets over the epochs are shown in Fig. A.6 and the values of internal proper motions with fitting errors are listed in Table A.2.

The measured internal proper motions are shown in Fig. 5. The amplitude of proper motions ranges from  $2.64 \pm 0.79 \text{ km s}^{-1}$  for clouddlet 7 to  $13.41 \pm 1.43 \text{ km s}^{-1}$  for clouddlet 8. Both are blueshifted and lying at the NW. On average the cloudlets have a velocity of  $5.45 \pm 3.21 \text{ km s}^{-1}$ . Most of the persistent cloudlets show motions nearly perpendicular to the plane of the edge-on disc-like structure, which best fits the 3D maser distribution and the time delays in the light curves of maser features as shown in Fig. 6. This possibly indicates that the methanol masers mainly trace the gas from a disc wind. However, the proper motions of the maser cloudlets in the SE part of the structure nearly align with the direction of the maser distribution (Fig. 4), indicating a flow along the structure, and may suggest a jet/outflow from the powering star. Somewhat random proper motions in this site may result from the method of proper motion studies we applied. We calculated the proper motions with respect to the barycentre of methanol distribution rather than the exciting star. We postpone a detailed analysis of plausible causes of the observed proper motions because the maser exhibits not only significant variations in the amplitude of flare and its profile, as the single dish monitoring proved, but also dramatic changes in the morphology at four cycles spanned only 1.6 yr. We note that only about one-fourth of cloudlets was identified as persistent. It is also possible that variations in the relative intensity of nearby maser spots with similar radial velocities at the observed epochs mimic



**Fig. 5.** Proper motions of the methanol maser cloudlets in G107 as measured relative to their centre of motion denoted by small black cross. Cones represent the proper motions in each cloudlet listed in Table A.2 and their colours and apertures correspond to the radial velocity (colour bar) and uncertainty, respectively. The velocity scale is shown at the top right corner. The dashed ellipse displays the model of ring derived from the 3D structure analysis shown in Fig. 6, where the yellow star shows the position of fitted centre of the ring. The grey star denotes the position of the 1.3 mm continuum peak (Palau et al. 2011).

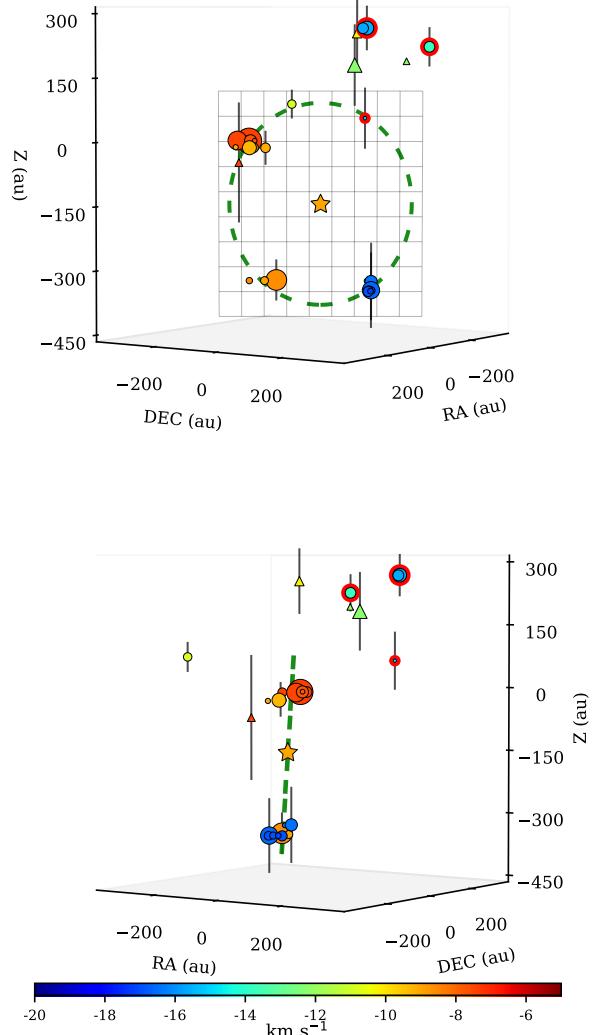
proper motion (Sugiyama et al. 2016). Further VLBI observations at high cadence could help interpret the proper motions in the target.

## 4. Discussion

### 4.1. Correlated variability of OH and $\text{CH}_3\text{OH}$ masers

Our observations definitively demonstrate that the intensity of the OH 1665 and 1667 MHz masers tightly follows that of the  $\text{CH}_3\text{OH}$  masers (Fig. 1) with a  $\lesssim 0.1P$  phase lag at 1665 MHz (Fig. 2). For the time intervals of simultaneous observations of both species the integrated flux densities of the 1665/1667 MHz and 6.7 GHz lines are well correlated (Fig. 7). This suggests similar excitation mechanisms for both molecules and that the masers come from the same gas volume.

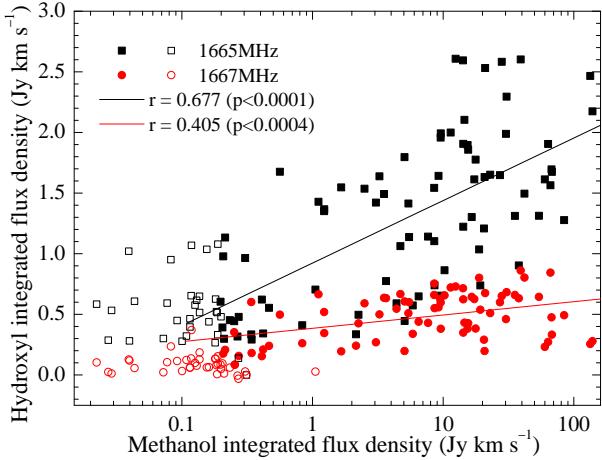
Figure 4 demonstrates that the OH maser cloudlet near  $-7.4 \text{ km s}^{-1}$  lies less than 56 mas (43 au) from the  $\text{CH}_3\text{OH}$  cloudlet near  $-9.1 \text{ km s}^{-1}$  (Table A.1). The OH cloudlet near  $-10.3 \text{ km s}^{-1}$  coincides within 57 mas (44 au) with the methanol cloudlet at  $-16.5 \text{ km s}^{-1}$ . In turn, the strongest and blue-shifted OH cloudlets at about  $-12.2 \text{ km s}^{-1}$  located  $\sim 250$  mas north of the region of persistent methanol masers (Fig. 4) coincide within  $< 120$  mas (100 au) with an area at which the methanol maser emission of velocity between  $-15.6$  and  $-13.7 \text{ km s}^{-1}$  appeared during the second epoch (MJD 57 444) of VLBI observations (Fig. A.4). Since the flux densities of both masers are strongly affected by spatial filtering by the VLBI, as Fig. A.4 depicts the above discussed spatial differences could be much smaller for low brightness extended emission. Thus, the hydroxyl and methanol masers originate in basically the same gas volume at a distance of 200–420 au from the central protostar. We note that the projected size of  $\sim 450$  au of the methanol maser region during the three epochs is close to the lower limit of the



**Fig. 6.** Three-dimensional maser structure of G107 obtained with the VLBI data and time delays of persistent feature peaks. The 6.7 GHz methanol and 1665 MHz hydroxyl cloudlets are shown as circles and triangles, respectively. The size of symbol is proportional to the logarithm of brightness and its colour corresponds to the LSR velocity scale shown in the horizontal wedge. The vertical bar for one cloudlet in each spectral feature denotes the uncertainty of the delay. The fitted position of protostar (orange star symbol) coincides within  $1\sigma$  with the measured position of the 1.3 mm continuum emission (Palau et al. 2011). The top and bottom panels show face-on and edge-on view of the ring-like structure. The green dashed circle in the top panel denotes the best fit to the methanol data at the fourth epoch of VLBI observation, whereas the green dashed line in the bottom panel indicates the plane of this ring-like structure. The symbols surrounded by the red circles represent the methanol maser cloudlets seen only at the fourth epoch of VLBI observations.

disc sizes observed in intermediate-mass embedded protostars (Beltrán & de Wit 2016). In such environments the hydroxyl and methanol masers may probe gas of different kinematic properties excited by the common central object. Previous surveys revealed that the majority (70–80%) of the 1.6 GHz hydroxyl masers around HMXSOs have associated 6.7 GHz methanol maser emission (Caswell 1998; Szymczak & Gérard 2004) with many sites exhibiting a close spatial coincidence.

The mechanisms of OH pumping were widely discussed (e.g. Guilloteau et al. 1981; Dewangan et al. 1987; Kylafis & Norman 1990; Cesaroni & Walmsley 1991; Cragg et al. 2002) and there is



**Fig. 7.** Correlation between the OH 1.665/1.667 GHz and  $\text{CH}_3\text{OH}$  6.7 GHz maser integrated flux densities. The open symbols denote the upper limits for at least one transition. The solid lines show the best fits to the data.

now a consensus that far-infrared dust emission and line overlap are essential to a feasible pumping scheme, which accounts for the observed OH maser sources. In G107 we note the dominance of 1665 MHz line over 1667 MHz by a factor of three, which is typical for HMXSOs (Caswell & Haynes 1987) and implies the effect of line overlap (Cragg et al. 2002). The VLBI observations revealed that the ratio of brightness temperatures for the  $\text{CH}_3\text{OH}$  6.7 GHz and OH 1665 MHz is 43 and 5.1 for the red- and blue-shifted cloudlets, respectively, implying a kinematic temperature of 110–120 K and dust temperature  $>250$  K according to the model by Cragg et al. (2002; their Figs. 5 and 6) with fixed values of the hydrogen density of  $10^7 \text{ cm}^{-3}$  and OH column density of  $10^{17} \text{ cm}^{-2}$ . Model calculations demonstrated that 6.7 GHz methanol masers share a broad range of physical parameters with 1.6 GHz hydroxyl masers (Cragg et al. 2002) and the coincidence of these transitions in G107 supports common excitation processes which produce the population inversions simultaneously in both molecules and result in synchronised variability of maser flux densities.

#### 4.2. Anti-correlated variability of $\text{CH}_3\text{OH}$ and $\text{H}_2\text{O}$ masers

The present study reinforces our previous finding of alternating variations of the methanol and water masers in the source (Szymczak et al. 2016). The methanol maser flare exactly coincides with the water maser minimum for all 33 cycles monitored over  $\sim 3.5$  yr (Figs. 1 and 2). Furthermore, the water maser emission that appeared outside the methanol velocity range, i.e. from  $-23$  to  $-17 \text{ km s}^{-1}$  and greater than  $-5 \text{ km s}^{-1}$ , also shows periodic variability. All the water maser cloudlets lie well within the methanol bounded region of projected size less than 200 au (Fig. 4). Thus, it is likely that the methanol and water masers operate in regions of similar physical parameters, where a factor causing periodic changes is efficient.

The excitation mechanisms of the 6.7 and 22 GHz masers are understood in general terms; the methanol line is likely inverted by infrared photons (e.g. Sobolev et al. 1997; Cragg et al. 2002) while the water line is predominantly collisionally pumped (e.g. Elitzur et al. 1989; Gray et al. 2016). Theoretical models generally indicate that the gas density and kinetic temperature favouring both maser transitions do not overlap (Cragg et al. 2002; Gray et al. 2016). However, modelling of the 22 GHz

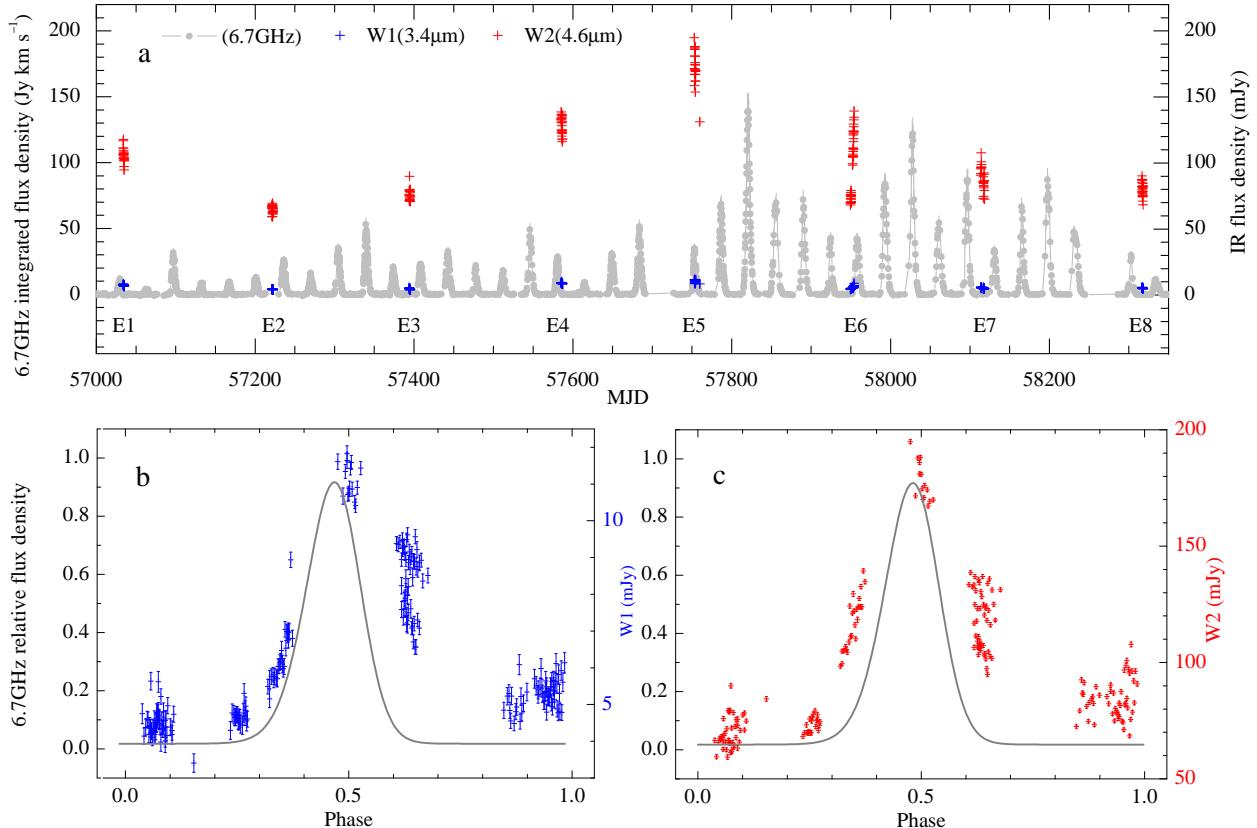
maser excitation with up-to-date molecular data (Nesterenok 2013) supports the view that the inversion is most efficient when the dust temperature is much lower than the gas temperature and an increase of the infrared emission reduces the maser gain (Yates et al. 1997).

The methanol flare profile in G107 is tightly related to the infrared intensity variations (Sect. 4.4) supporting the radiative pumping models (Cragg et al. 2002). In response to a luminosity burst, traced by the methanol maser, the dust temperature reaches a maximum just at the methanol maser peak. This event reduces the ratio of the dust to kinetic temperature causing a decrease in the water maser intensity. As the heating and cooling times of dust grain are much shorter than those of gas (e.g. Johnstone et al. 2013) then we can expect the decaying branch of the  $\text{H}_2\text{O}$  light curve to be much steeper than the rising branch. This is exactly what we observed for the velocity-integrated emission (Fig. 2) and  $R_{rd}$  ratio of the maser features (Table 2).

#### 4.3. Three-dimensional model from VLBI maps and time delays

In order to find the 3D methanol maser structure we applied the same procedure as in Olech et al. (2019). The fourth epoch of VLBI data was used as the reference of the typical emission structure. The blue-shifted methanol emission seen only at the second VLBI epoch (Fig. A.4) and the OH emission mapped with VLBA were added to include all the methanol and hydroxyl features with measured time delays (Table 2). We assume that the periodic variations of the masers are caused by changes in the stellar and accretion luminosity of the central object which affect the infrared pump rate, while the time delays are due to differences in the photon paths between the cloudlets seen by the distant observer. We note that each feature in the spectrum was a superposition of the emission from spatially separated groups of cloudlets. Both the position of the central object and the line-of-sight distances between the cloudlets were fitted to the observed time delays.

The recovered structure is presented in Fig 6. All persistent methanol features, with the exception of the  $-11.03 \text{ km s}^{-1}$  feature, lie in a plane inclined by  $86 \pm 16^\circ$  to the plane of the sky that is equidistant from the fitted position of powering object. The groups of cloudlets form a ring-line structure of radius  $242 \pm 5.8$  au and an average thickness of  $30.1 \pm 37.4$  au. This clearly indicates that the 6.7 GHz maser emission comes from a disc seen almost edge-on. The  $-11.03 \text{ km s}^{-1}$  methanol emission lies at a distance of 360 au from the protostar and above the fitted plane of the disc. The blue-shifted cloudlets detected at the second epoch of VLBI observations are 500–560 au from the central protostar. The erratic appearance of these cloudlets in VLBI observations and low relative amplitudes of variability seen in the single-dish monitoring are possibly related to a relatively low luminosity of the central object. We can expect that flares of the central star at only some cycles could be strong enough to heat up the dust in those distant regions of the circumstellar envelope. In this model the red-shifted OH emission ( $-7.4 \text{ km s}^{-1}$ ) lies close to the methanol red-shifted cloudlets, while that at blue-shifted velocities coexists with the erratic (seen during the second VLBI observation) methanol cloudlets at about 450 au from the central star. We conclude that assuming radiative pumping of the  $\text{CH}_3\text{OH}$  and OH masers, the structure of maser emissions and time delays of the flares of individual features observed in G107 can be satisfactorily explained by the disc model.



**Fig. 8.** Panel *a*: comparison of the methanol maser time series and the NEOWISE photometric data during this period. The epochs of IR observations are numbered with E1–E8. Panels *b* and *c*: NEOWISE light curves folded by the 34.4 d period for bands *W1* and *W2*, respectively. The grey line shows the phased and normalised methanol light curve as taken from the top panel of Fig. 2.

#### 4.4. Relationship between the methanol maser and IR variability

To explore the relationship between the maser flares and infrared emission in G107 we extracted the 3.4  $\mu\text{m}$  (*W1*) and 4.6  $\mu\text{m}$  (*W2*) time series from NEOWISE photometry. Details about the instrument and observations are given in Mainzer et al. (2011, 2014) and a guide for users of the NEOWISE data is presented by Cutri et al. (2015). The NEOWISE-R Single Exposure Source Table<sup>1</sup> was queried using a 5'' search radius and the data covering the maser monitoring period were found for eight observation cycles 1.3–2.0 d in length (hereafter called epochs). During each epoch, some 20–30 points were available with intervals of 0.066–0.262 d. The data with the frame quality score *qual\_frame* = 0 were rejected and 211 exposures of high quality (*ph\_qual* = “A”) were used. Their median signal-to-noise ratios at bands *W1* and *W2* are 47.6 and 83.4, respectively. The magnitudes of individual measurements were converted to fluxes following the WISE Explanatory Supplement<sup>2</sup>.

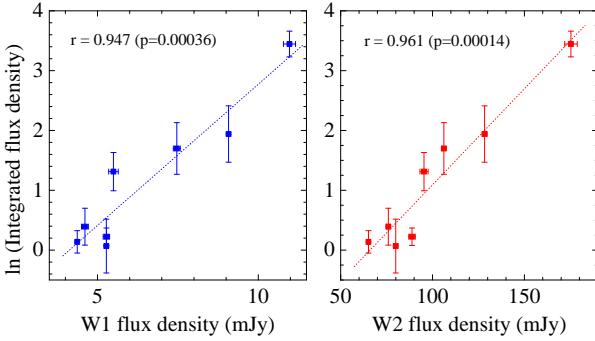
Figure 8a shows the time series of the 6.7 GHz integrated flux density and 3.4 and 4.6  $\mu\text{m}$  flux densities. It is clear that the NEOWISE observations at the second (E2, ~MJD 57 222), third (E3, ~MJD 57 395), seventh (E7, ~MJD 58 116), and eighth (E8, ~MJD 58 317) epochs were taken during a quiescent state, when the maser flux density was below our detection limit. For the remaining four epochs the IR observations were contemporaneous with active phases of the maser. We note

that the amplitude of the maser flares closest to the epochs of NEOWISE observations ranged from 10 to 40 Jy  $\text{km s}^{-1}$ , while over the whole monitoring period (Fig. 1) it varied from 3 to 150 Jy  $\text{km s}^{-1}$ . This may suggest that a long-term variability does not strongly affect the maser–infrared relationship at these eight epochs.

The phased infrared light curves are shown in Figs. 8b and c. There is a general agreement between the flare profiles at the infrared and maser wavelengths with the exception of a bump in the IR light curves at phase ~0.63. The NEOWISE data at this phase come from epochs E1 and E4 (Fig. 1a), when the maser peaks during the preceding activity cycles differed by a factor of 2.9. The ratio of average flux densities at E4 and E1 epochs is 1.21 for both IR bands. If we assume that this change in the IR fluxes makes a difference in the maser optical depth  $\Delta\tau = 1.21$  then for unsaturated maser regime we obtain  $e^{(\Delta\tau)} = 3.3$ , which is surprisingly consistent with the observed ratio of the maser peak flux densities of 2.9. This appears consistent with a scenario of pumping by infrared emission from dust surrounding the central star (Sobolev et al. 1997; Cragg et al. 2002). Stecklum et al. (2018) have established the light curve of G107 from ALLWISE and NEOWISE data folded by a 34.6 d maser period. The saw tooth shape of their curve differs from that presented here, since they did not have precise information about the phase of the maser flares. Figure 9 confirms the strong correlation between the 6.7 GHz line and IR flux densities and supports a radiative pumping of this unsaturated maser emission. Our results could be further refined using high cadence data at longer IR wavelengths to investigate in more details the pumping mechanisms.

<sup>1</sup> Available at [irsa.ipac.caltech.edu](http://irsa.ipac.caltech.edu)

<sup>2</sup> [wise2.ipac.caltech.edu/docs/release/neowise/expsup/](http://wise2.ipac.caltech.edu/docs/release/neowise/expsup/)



**Fig. 9.** Natural logarithm of the methanol integrated flux density measured in  $\text{Jy km s}^{-1}$  vs. band W1 (blue) and W2 (red) flux densities. The dotted lines denote least-squares fits to the data.

## 5. Summary and conclusions

The IMYSO G107 was observed at high angular resolution and monitored over the period 2015–2018 in the four maser transitions: 22.2 GHz  $\text{H}_2\text{O}$ , 6.7 GHz  $\text{CH}_3\text{OH}$ , 1665 and 1667 MHz OH. The main findings and conclusions are summarised as follows:

1. The periodic variations of the OH and  $\text{CH}_3\text{OH}$  masers are simultaneous to within less than 3 d, whereas those of the  $\text{H}_2\text{O}$  maser are delayed exactly by half of the period of 34.4 d, i.e. the water maser emission is significantly dimmed or even disappears at regular intervals just when the OH and  $\text{CH}_3\text{OH}$  masers peak.
2. We found a strong correlation between the 6.7 GHz integrated flux density and 3.4 and 4.5  $\mu\text{m}$  flux densities observed with the NEOWISE at eight epochs. This strongly supports theoretical models of radiative pumping of this maser transition.
3. VLBI observations confirmed the elongated distribution of the methanol emission with a clear velocity gradient. The 3D structure recovered from these data and time delay measurements of persistent features implies a ring-like structure of radius 240 au and thickness 30 au, while erratic methanol emission seen at one epoch is located 500 au from the exciting star. Thus, the observations are well explained with a disc model in which both OH and  $\text{CH}_3\text{OH}$  masers are radiatively pumped.
4. Measurements of internal proper motions of the methanol maser cloudlets indicate velocity vectors mostly orientated perpendicularly to the main axis of structure suggesting the presence of a disc wind.

Overall, the study provides strong support for models of radiative pumping of the OH and  $\text{CH}_3\text{OH}$  molecules in a working region between the disc and outflow. Future VLBI imaging will help constrain the kinematics of cloudlets of all the maser transitions and quasi-simultaneous observations in other radio and infrared bands, and allow us to better understand the complicated structure seen in the thermal lines (e.g. Palau et al. 2013).

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## Appendix A: Supplementary tables and figures

**Table A.1.** Parameters of the cloudlets during the three maser transitions observed over a timespan of 1.1 yr.

$\Delta\text{RA}$ (mas)	$\Delta\text{Dec}$ (mas)	$V_p$ (km s $^{-1}$ )	$FWHM$ (km s $^{-1}$ )	$S_p$ (Jy beam $^{-1}$ )	$S_f$ (Jy beam $^{-1}$ )
<i>OH 1665 MHz (VLBA)</i>					
-14.954	-294.106	-7.42	0.146	0.310	0.318
-119.285	135.742	-10.27	0.892	0.525	0.484
-65.271	373.965	-12.12	—	0.22	—
61.145	247.472	-12.22	0.314	1.461	1.547
<i>OH 1665 MHz (EVN)</i>					
-9.774	-300.027	-7.56	—	0.096	—
-115.510	130.847	-10.40	1.130	0.485	0.479
-49.913	372.111	-12.00	—	0.050	—
67.252	241.243	-12.33	0.567	0.744	0.773
<i>H<sub>2</sub>O 22.2 GHz</i>					
-13.313	-74.900	-7.58	1.186	21.543	12.968
-204.755	41.600	-15.97	1.048	29.513	30.837
-202.988	36.955	-16.20	0.813	32.567	32.353
-205.331	40.530	-16.60	0.804	82.596	82.593
-197.815	39.607	-16.71	—	3.789	—
-181.082	7.593	-16.90	1.331	18.218	6.635
-205.179	39.908	-17.30	—	27.191	—
-195.511	30.094	-17.31	0.705	41.828	40.240
-179.586	5.135	-17.44	0.790	42.912	40.588
-179.410	3.670	-19.17	0.559	40.56	40.385
<i>CH<sub>3</sub>OH 6.7 GHz epoch 4 (MJD 57 681)</i>					
99.886	-174.277	-7.13	—	8.865	—
88.092	-167.362	-7.21	0.287	1.741	1.656
97.987	-167.545	-7.36	0.332	15.630	16.429
117.979	-159.029	-7.39	0.246	0.652	0.674
65.320	-226.854	-7.40	—	0.434	—
109.386	-209.002	-7.40	0.275	3.782	3.891
100.701	-159.262	-7.73	0.295	0.969	0.939
92.954	-147.215	-7.85	—	0.105	—
82.251	-176.911	-8.06	0.442	0.101	0.101
65.986	-190.234	-8.55	0.311	0.189	0.200
-7.284	-127.389	-8.59	0.280	5.807	5.829
35.387	-147.094	-8.99	0.287	0.338	0.329
50.225	-206.324	-9.01	0.292	0.310	0.320
36.559	-270.120	-9.08	—	0.129	—
10.817	-160.301	-9.12	0.122	0.515	0.517
30.766	-198.142	-9.20	0.211	0.569	0.568
38.960	-209.079	-9.26	0.228	1.389	1.420
-302.581	-269.756	-10.91	0.215	0.328	0.382
-289.855	103.608	-14.95	—	0.082	—
-181.412	134.757	-16.33	0.598	0.112	0.114
-171.342	159.929	-16.47	0.749	1.102	0.956
-236.113	116.447	-16.58	0.510	2.974	2.859
-234.022	109.275	-16.87	—	0.412	—
-225.448	125.518	-16.90	0.364	0.230	0.262
-192.541	133.441	-16.91	0.285	0.595	0.592
-208.412	132.691	-17.11	0.361	0.165	0.154

**Notes.** Coordinates ( $\Delta\text{RA}$ ,  $\Delta\text{Dec}$ ) are relative to the map origin in Fig. 4. Most of cloudlets are parametrized with the fitted brightness ( $S_f$ ) and line width at half maximum (FWHM). The 6.7 GHz maser cloudlets seen at epoch 2 are added to show their different appearance relative to that at the three other epochs.

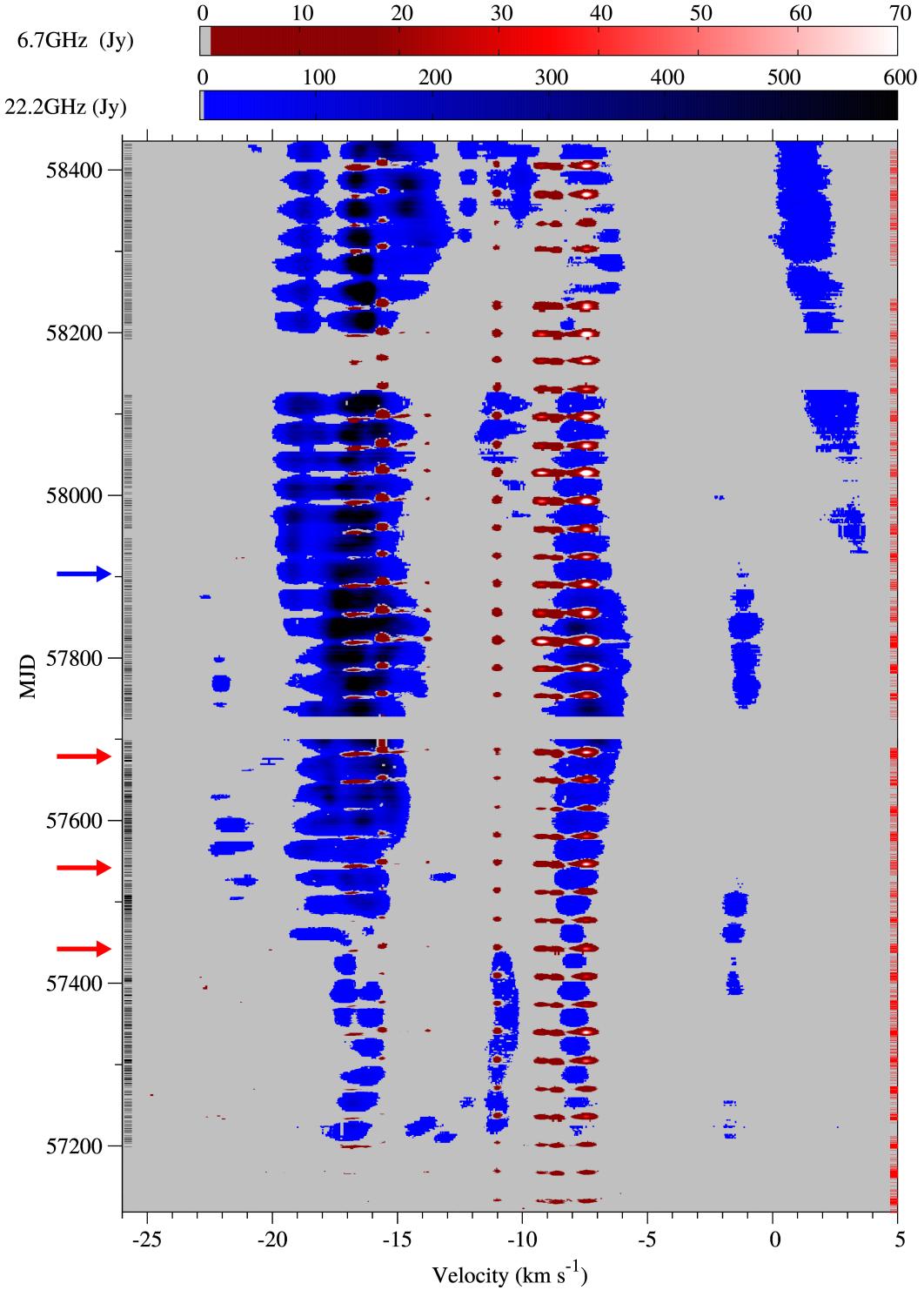
**Table A.1.** continued.

$\Delta\text{RA}$ (mas)	$\Delta\text{Dec}$ (mas)	$V_p$ (km s $^{-1}$ )	$FWHM$ (km s $^{-1}$ )	$S_p$ (Jy beam $^{-1}$ )	$S_f$ (Jy beam $^{-1}$ )
<i>CH<sub>3</sub>OH 6.7 GHz epoch 2 (MJD 57 444)</i>					
89.126	-168.899	-7.05	0.298	0.370	0.384
102.004	-173.640	-7.12	0.294	7.132	7.243
111.537	-153.678	-7.16	0.252	0.197	0.206
116.095	-206.232	-7.31	0.311	0.599	0.629
100.052	-167.717	-7.35	0.367	11.853	12.159
105.166	-211.788	-7.50	0.220	1.094	1.113
102.335	-159.127	-7.79	—	1.170	—
95.592	-146.682	-7.85	—	0.052	—
84.405	-176.262	-8.04	0.337	0.114	0.117
95.083	-162.762	-8.19	—	0.057	—
85.779	-159.710	-8.50	0.343	0.058	0.058
67.988	-189.282	-8.54	0.262	0.155	0.156
-6.336	-126.889	-8.59	0.293	1.195	1.197
27.436	-236.224	-8.92	—	0.107	—
51.579	-207.241	-9.01	0.202	0.313	0.309
39.190	-270.110	-9.08	—	0.115	—
32.588	-199.424	-9.25	—	0.077	—
40.285	-208.908	-9.28	0.223	0.669	0.650
304.259	-261.180	-10.96	0.227	1.408	1.427
-107.812	438.657	-13.70	0.223	0.645	0.670
144.067	345.034	-14.81	—	0.054	—
154.643	359.164	-15.50	0.267	1.269	1.330
157.612	345.580	-15.63	0.238	0.707	0.709

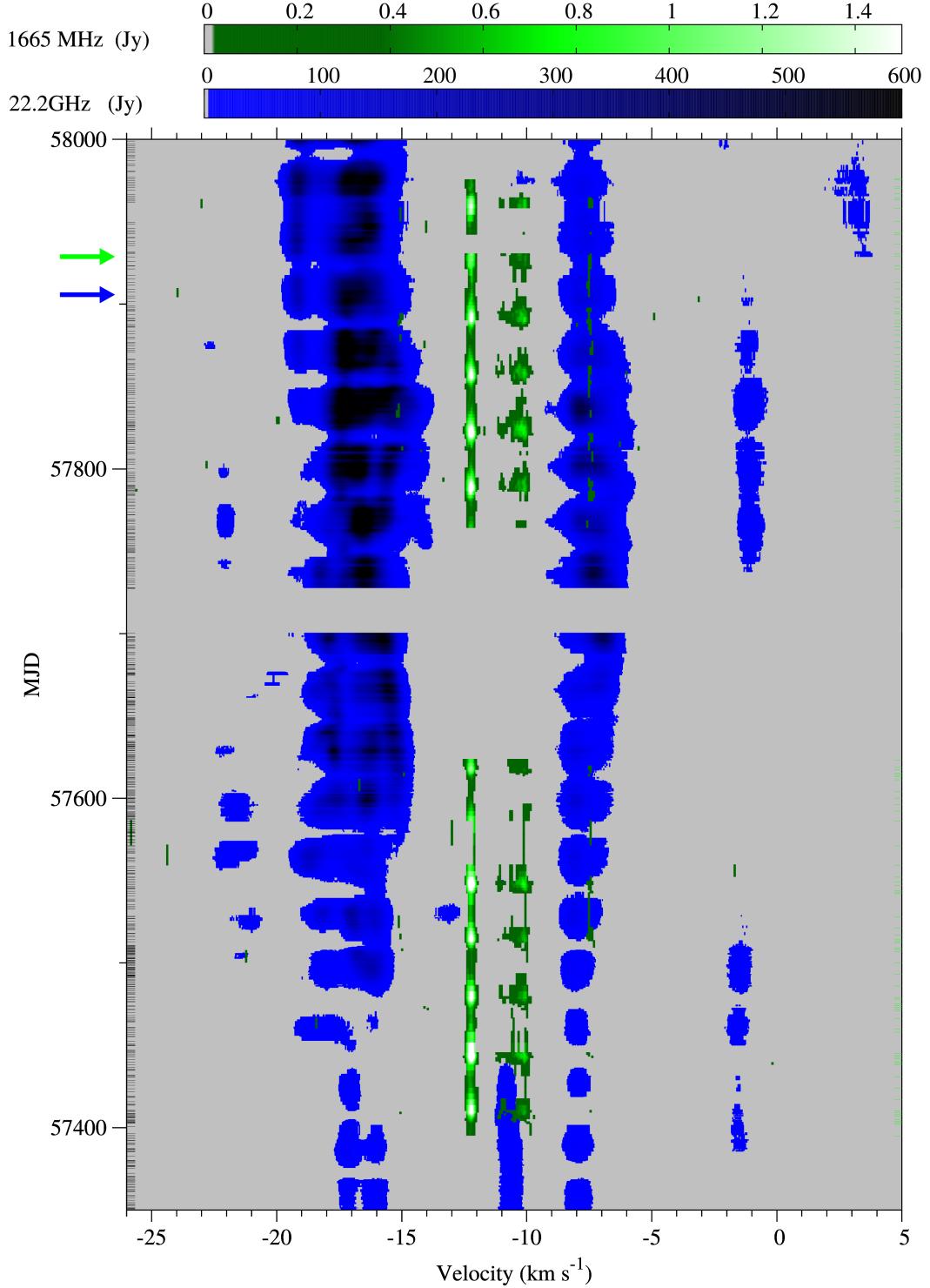
**Table A.2.** Proper motions of the 6.7 GHz maser cloudlets in G107.

Cloudlet	$\Delta\text{RA}$ (mas)	$\Delta\text{Dec}$ (mas)	$V_p$ (km s $^{-1}$ )	$S_p$ (Jy beam $^{-1}$ )	$S_f$ (Jy beam $^{-1}$ )	$V_{\text{RA}}$ (km s $^{-1}$ )	$V_{\text{Dec}}$ (km s $^{-1}$ )
1	90.409	-165.812	-7.12	0.119	0.121	$4.14 \pm 0.70$	$0.40 \pm 0.72$
2	101.875	-164.401	-7.42	7.575	8.181	$-0.19 \pm 0.76$	$-3.98 \pm 0.40$
3	85.773	-173.574	-8.09	0.053	0.056	$1.28 \pm 0.46$	$-4.36 \pm 0.55$
4	41.593	-268.280	-9.07	0.048	0.052	$-3.53 \pm 0.04$	$-0.74 \pm 1.02$
5	35.316	-197.933	-9.25	0.187	0.196	$-2.07 \pm 0.09$	$6.64 \pm 1.62$
6	43.614	-206.295	-9.28	0.328	0.321	$-3.69 \pm 2.60$	$-2.14 \pm 0.89$
7	-231.936	118.904	-16.58	0.016	—	$-1.26 \pm 0.74$	$-2.32 \pm 0.28$
8	-168.709	156.797	-16.68	0.032	0.033	$3.17 \pm 0.55$	$13.03 \pm 1.32$

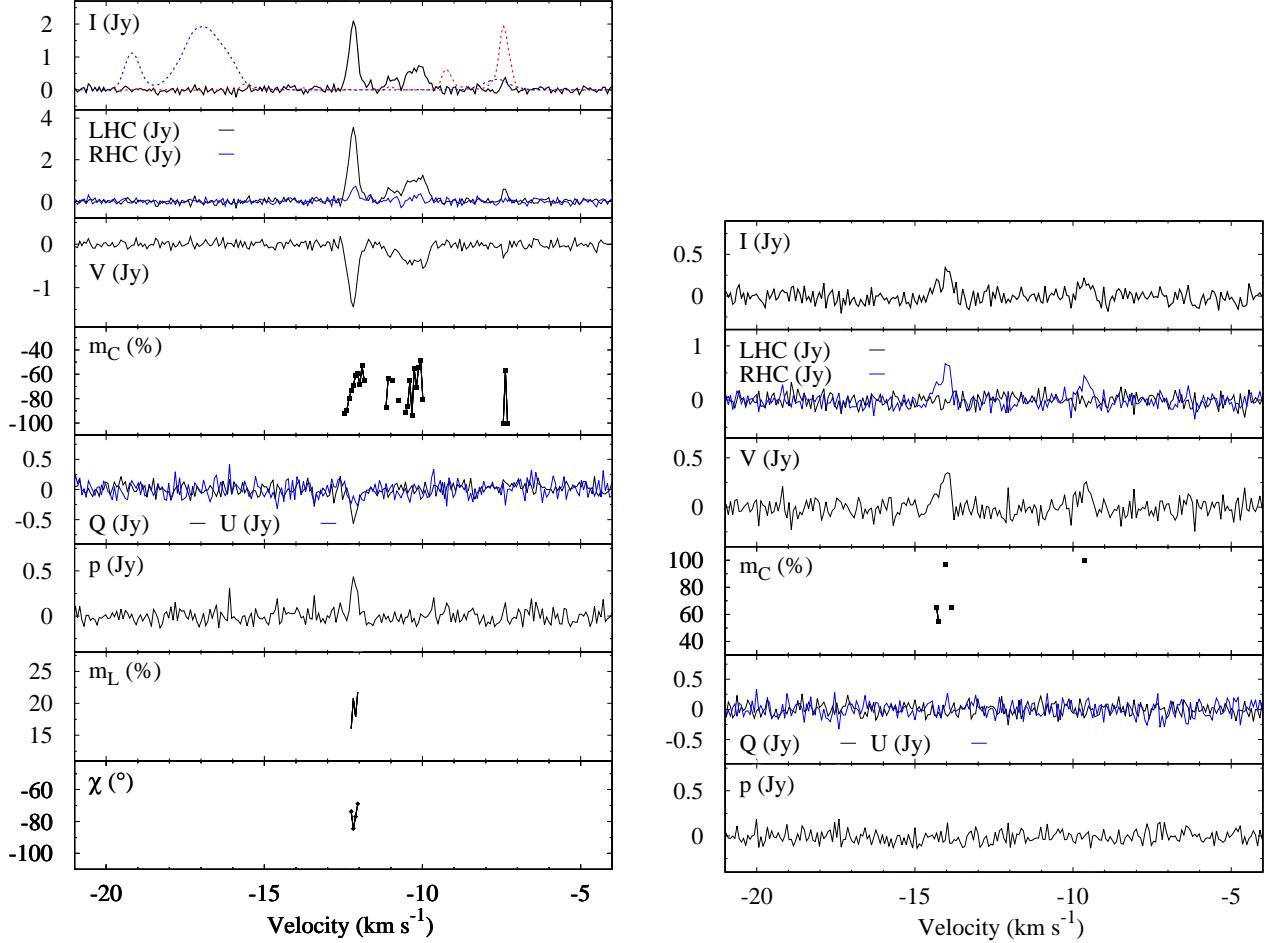
**Notes.** The values of internal proper motions of the cloudlets and their uncertainties are listed in Cols. 7 and 8. Designations of Cols. 2–6 are identical to those in Table A.1.



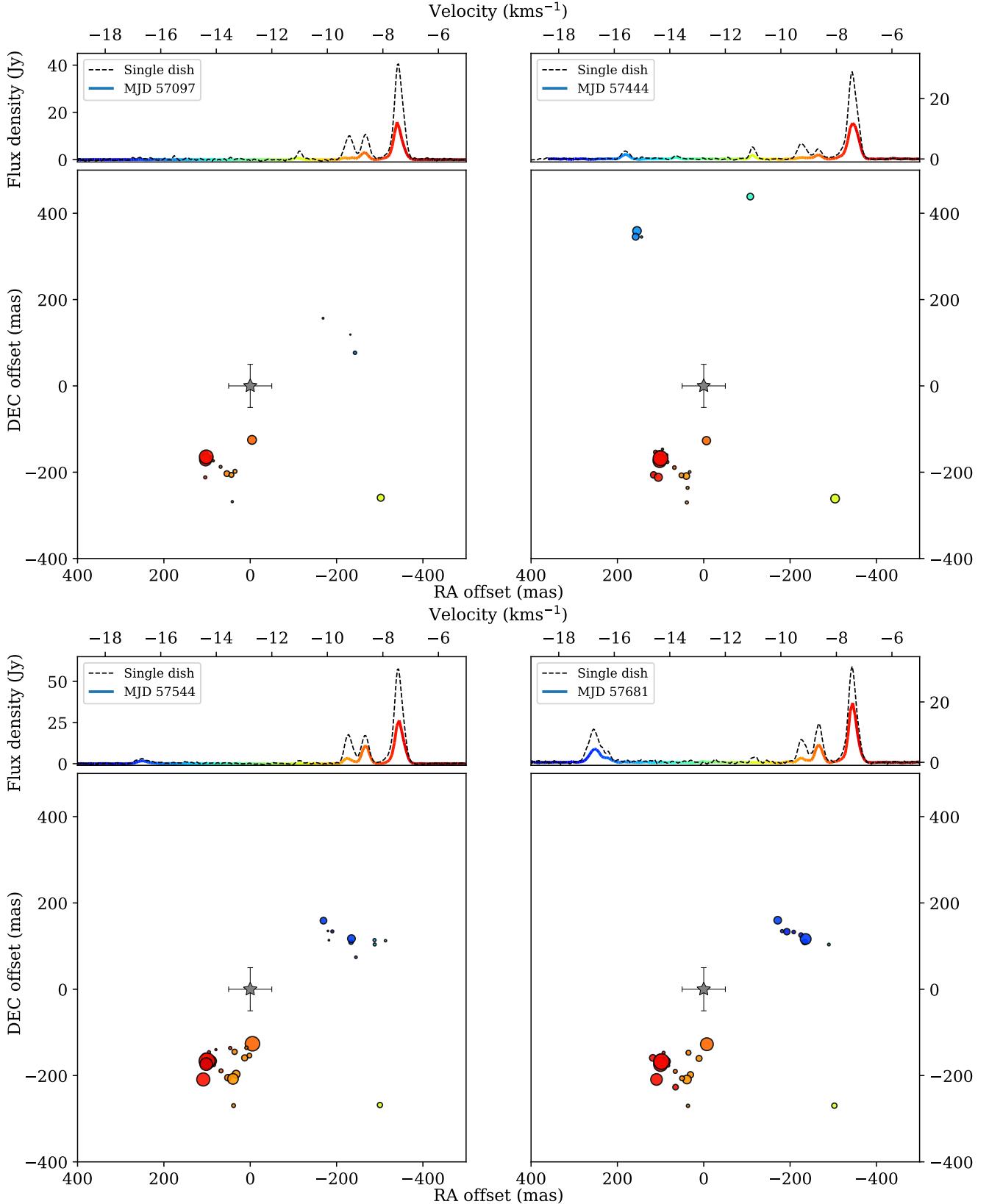
**Fig. A.1.** Dynamic spectra of the methanol and water maser lines of G107. Radial velocity is measured with respect to the LSR. The horizontal bars in the left (black) and right (red) ordinates correspond to the dates of the observed spectra of water and methanol lines, respectively. The red and blue arrows indicate the epochs of VLBI observations of methanol and water, respectively.



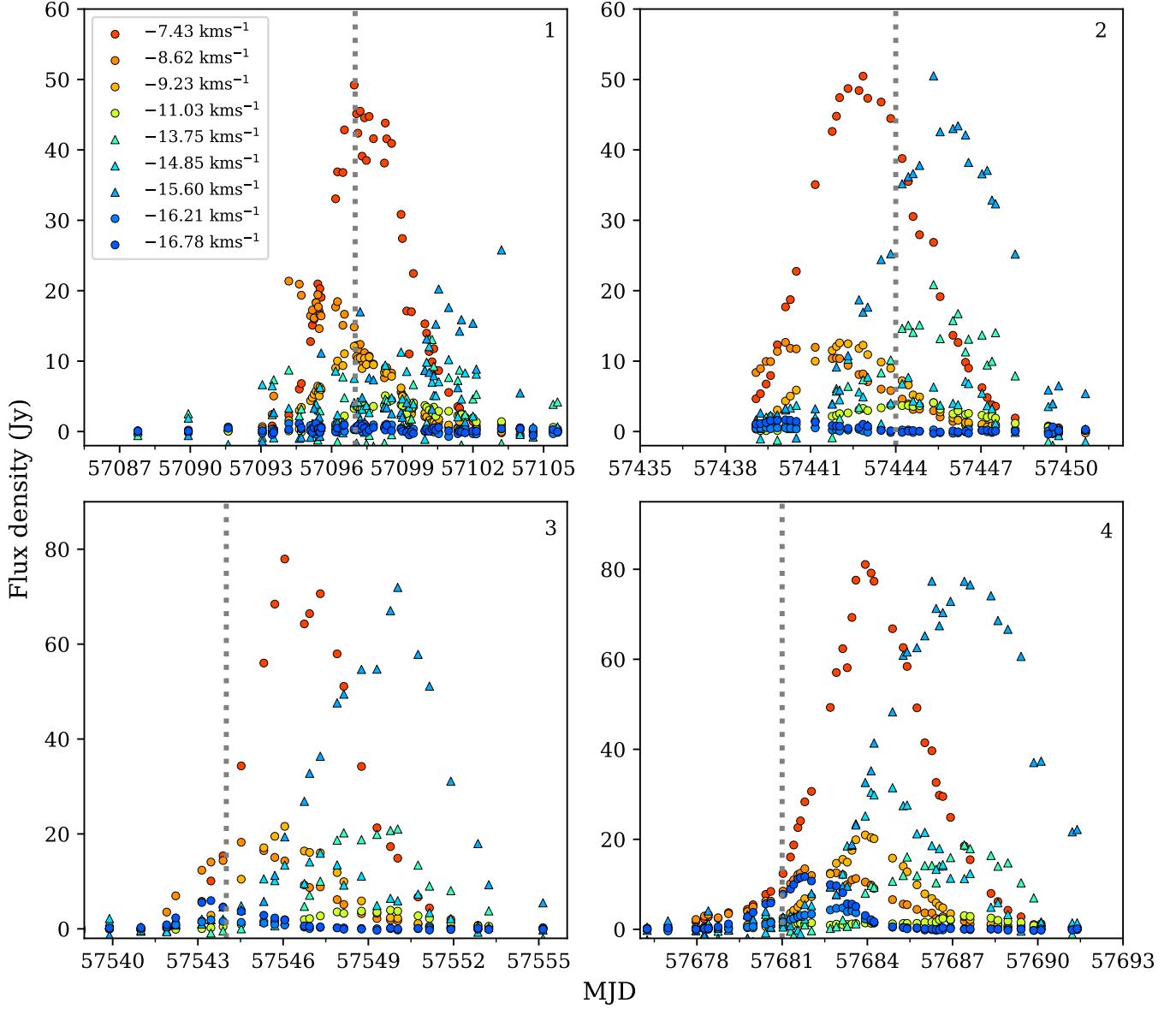
**Fig. A.2.** Dynamic spectra of the water and hydroxyl maser lines of G107. Radial velocity is measured with respect to the LSR. The horizontal bars in the left (black) and right (green) ordinates correspond to the dates of the observed spectra of water and hydroxyl lines, respectively. The blue and green arrows indicate the epochs of VLBI observations of water and hydroxyl, respectively.



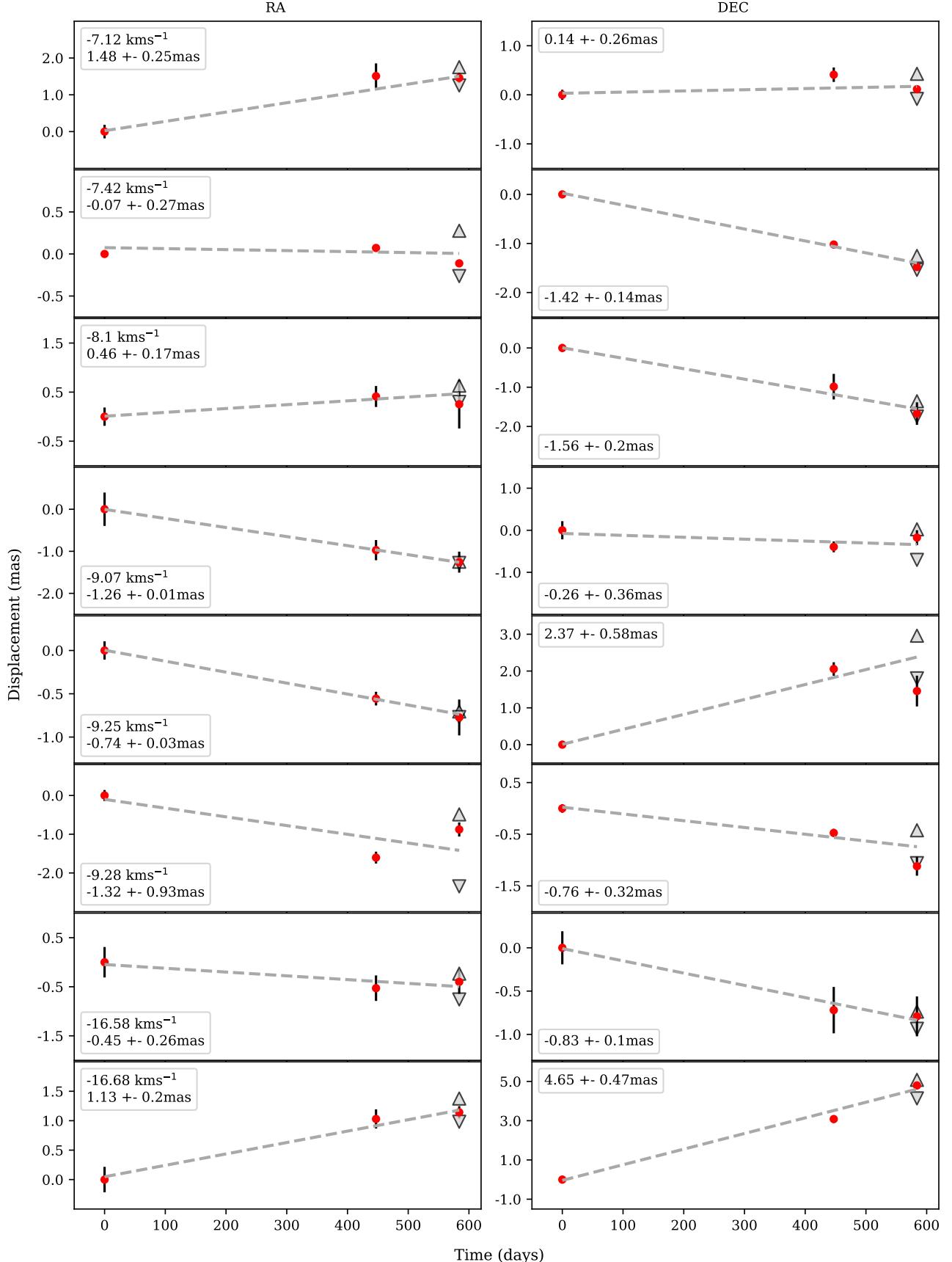
**Fig. A.3.** Full polarisation spectra of the OH 1665 MHz (*left*) and 1667 MHz (*right*) maser lines of G107 taken on March 9, 2017 (MJD 57 821). Spectra of the Stokes  $I$ , LHC, RHC, Stokes  $V$ , degree of circular polarisation ( $m_C$ ), Stokes  $Q$ ,  $U$ , linearly polarised flux density ( $p$ ), degree of linear polarisation ( $m_L$ ), and polarisation position angle  $\chi$  are shown from top to bottom. In the *top left panel* the Stokes  $I$  spectra of  $\text{CH}_3\text{OH}$  (red dotted line) and  $\text{H}_2\text{O}$  (blue dotted line) maser lines divided by a factor of 50 and 200, respectively, are added.



**Fig. A.4.** 6.7 GHz maser spectra and spatial structures during all epochs of EVN observations. Coloured lines show the cross-correlated VLBI spectra and colour corresponds to velocity of maser cloudlets shown in the bottom panels. The emission is filtered out by the EVN beam by more than 40–60% for most of the spectral features. The size of each circle is scaled with the logarithm of peak flux density. The symbol of grey star with the error bars indicates the position of the 1.3 mm continuum emission peak (Palau et al. 2013).



**Fig. A.5.** Light curves of 6.7 GHz maser features around the EVN observations. The dotted lines indicates the dates of the EVN experiments numbered in the top right corner. The features in which emission was detected only at the second epoch are denoted by triangles and their flux density was scaled up by a factor of 15. Differences in the maser morphology seen by EVN (Fig. A.4) are clearly related to the flare phase.



**Fig. A.6.** Proper motion measurements vs. time for each cloudlet with RA and Dec components on the left and right panels, respectively. Red points show measured values with respective errors. Dashed lines show best fits of linear functions. Top and bottom triangles show the error range of the fit. The proper motion with estimated errors for each axis is shown in boxes. The velocity of the component concerned is given in the *left panel*. The final result is shown in Fig. 5.